

THE OPTICAL SPECTRUM OF THE SN 1006 SUPERNOVA REMNANT REVISITED

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ABSTRACT

We present the deepest optical spectrum acquired to date of Balmer-dominated shocks in the NW rim of SN 1006. We detect the broad and narrow components of H α , H β and H γ and report the first detection of the He I λ 6678 emission line in this supernova remnant. We may have detected, at the 1.5σ level, faint He II λ 4686 emission. We measure a full width half maximum of 2290 ± 80 km s⁻¹ in the broad component H α line, with broad-to-narrow flux ratios of $0.84^{+0.03}_{-0.01}$ and $0.93^{+0.18}_{-0.16}$ in H α and H β , respectively. To match these observations, our nonradiative shock models require a low degree of electron-proton equilibration at the shock front, $T_e/T_p \leq 0.07$, and a shock speed of 2890 ± 100 km s⁻¹. These results agree well with an earlier analysis of ultraviolet lines from SN 1006. The He I/H α and He I/He II flux ratios also indicate low equilibration. Furthermore, our models match the observations for mostly ionized ($\sim 90\%$) preshock H and mostly neutral ($\gtrsim 70\%$) preshock He, respectively. We conclude that the high H ionization fraction cannot be explained by either photoionization from the reverse shock or relic ionization from EUV photons released in the 1006 A.D. supernova. The most plausible explanation appears to be photoionization from the Galactic Lyman continuum.

Subject headings: ISM: supernova remnants: individual (SN 1006)–ISM: kinematics and dynamics, shock waves

1. INTRODUCTION

Among the known galactic supernova remnants (SNRs), SN 1006 stands out as one of the few with low extinction ($E(B - V) = 0.11$, Schweizer & Middleditch 1980), a well determined age and a reasonably well known distance ($D \sim 2$ kpc, Kirshner, Winkler & Chevalier 1987 (hereafter KWC87), Long *et al.* 1988, Laming *et al.* 1996, Winkler & Long 1997). This SNR was probably produced by a Type Ia explosion (Schaefer 1996) and appears as a limb brightened shell $30'$ in diameter at X-ray (Willingale *et al.* 1996, Winkler & Long 1997) and radio (Reynolds & Gilmore 1986) wavelengths. Using the Hopkins Ultraviolet Telescope (HUT), Raymond *et al.* (1995) first detected ultraviolet emission lines of He II λ 1640, C IV λ 1548, 1551, N V λ 1239, 1243 and O VI λ 1032, 1038 in SN 1006. Raymond *et al.* (1995) observed equal velocity widths for the emission lines detected by HUT. Combining this information with the relative strengths of the He II, C IV, N V and O VI lines, Laming *et al.* (1996) concluded that the lines observed by HUT were due to collisional excitation in 2600 km s⁻¹ nonradiative shocks with little or no electron-ion temperature equilibration at the shock front. In this case the high shock speed and low density ($n \lesssim 1$ cm⁻³) result in negligible radiative cooling downstream (hence the term ‘nonradiative shock’). All of the line emission is produced in a narrow ionization zone close to the shock front.

In the optical, SN 1006 appears as a faint network of

Balmer-dominated filaments (van den Bergh 1976, Long, Blair & van den Bergh 1988, Winkler & Long 1997). These filaments emit an optical spectrum dominated by the Balmer lines of H, produced in the same ionization zone as the ultraviolet lines. Balmer line profiles consist of two components: a narrow component produced when cold ambient neutrals are overrun by the shock and a broad component produced when charge exchange with protons generates a hot neutral population behind the shock (Chevalier & Raymond 1978, Chevalier, Kirshner & Raymond 1980). The broad component width is directly proportional to the postshock proton temperature, while the ratio of broad to narrow flux is sensitive to the electron-proton temperature equilibration and the preshock neutral fraction. These diagnostics are invaluable tools for measuring shock speeds in supernova remnants and studying the efficiency of collisionless heating processes in high Mach number, nonradiative shocks.

2. SPECTROSCOPIC OBSERVATIONS

We obtained a deep long-slit spectrum of the NW filament in SN 1006 from the 4-m Blanco telescope at CTIO on 1998 June 24 (UT). The RC Spectrograph was used with the 527 line mm⁻¹ grating KPGL3, the Blue Air Schmidt camera, and the Loral 3K CCD. This combination covered the spectral range 3760 – 7480 Å at a dispersion of 1.22 Å pixel⁻¹. For our slit width of $2''0$, the instrumental resolution was approximately 4.5 Å (205 km s⁻¹ at H α) along the entire spectrum. The plate scale for this

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configuration was $0''.51 \text{ pixel}^{-1}$ and the slit length was $5''$. We rotated the slit to $\text{PA} = 63.9^\circ$ and placed its center at coordinates $15^{\text{h}}02^{\text{m}}13^{\text{s}}.5, -41^\circ45'22''$ (2000), as shown in Fig. 1. We chose this position to cut obliquely across the NW filament in order to capture as much of the bright region as possible while giving enough background sky, away from the filament, for a good sky subtraction. It is virtually identical to the position used by KWC87. We obtained 4 exposures at this location, with a total exposure time of 8400 s, which we combined into a single two-dimensional spectrum following initial reduction.

The data were reduced using standard IRAF⁶ procedures of bias subtraction, flat fielding and illumination correction. Spectra of a HeNeAr lamp taken at the beginning and end of the series of object frames were used for wavelength calibration. We subtracted the night sky contribution from the two-dimensional spectrum using emission from the two ends of the slit. The $\text{H}\alpha$ line profiles (both broad and narrow components) were nearly constant along the observed filament; therefore, we integrated the emission along the slit to obtain a single high signal-to-noise spectrum from a $51''$ section of the filament (Figs. 2 and 3). Exposures of several spectrophotometric standard stars from Hamuy et al. (1994) were used for flux calibration. We estimate an absolute photometric accuracy of 20% for our quoted emission line fluxes.

3. DETECTED EMISSION LINES

Among the most interesting new features in our optical data are the detection of broad $\text{H}\beta$ and $\text{He I } \lambda 6678$. To our knowledge, this is the first detection of the $\lambda 6678$ line in a pure nonradiative shock. The neutral He atoms passing downstream are unaffected by the electromagnetic turbulence at the shock front. Therefore, neutral He atoms remain cool throughout the shock and produce narrow emission lines unresolved in our data. On the other hand since He^+ is an ion, it is heated at the shock front by the same collisionless processes which heat the electrons and ions. Therefore, we expect the $\text{He II } \lambda 4686$ line to be broad and more difficult to detect than the $\text{He I } \lambda 6678$ line. There is a hint of the $\text{He II } \lambda 4686$ line at a low (1.5σ) statistical significance. The surface brightnesses of the detected emission lines from NW SN 1006 appear in Table 1. We have used $E(B - V) = 0.11$ (Schweizer & Middleditch 1980) to obtain dereddened line intensities from our measured spectra.

4. SPECTRAL LINE FITS

We fit the line profiles using a standard nonlinear χ^2 reduction routine. We assumed Gaussian profile shapes for all lines and fitted the $\text{H}\alpha$, He I , $\text{H}\beta$, $\text{H}\gamma$ and He II lines separately. Each Balmer line profile was fit with two Gaussians, one each for the narrow and broad components. The He lines were fit with single Gaussians. To reduce the required degrees of freedom, we set the widths of the narrow component Balmer lines and He I line equal to the instrumental resolution during the fitting process.

Due to the faintness of the $\lambda 4686$ emission, we cannot estimate the width of this line from the fits themselves. Therefore, we set the velocity width of the He II emission

line width equal to that of the $\text{H}\alpha$ line. This assumption is supported by the HUT analysis of Raymond *et al.* (1995), who found that to within the errors, the velocity widths of the $\text{He II } \lambda 1640$ and broad Ly β lines are equal (i.e., electron-ion and ion-ion temperature equilibration is inefficient in SN 1006). The best fit for the $\text{He II } \lambda 4686$ line profile appears in Fig. 4. We note that due to the large width and faintness of the He II line, the flux predicted by the line fit is quite sensitive to the baseline level: the lower the baseline, the higher the predicted He II flux. Both the statistical uncertainty and the systematic baseline uncertainty contribute to the large error bars on the flux. For the purpose of our subsequent analysis, the most important quantity here is the upper limit on the He II line flux.

Our best fit for the $\text{H}\alpha$ broad component width is $2290 \pm 80 \text{ km s}^{-1}$, agreeing very well with the value of $2310 \pm 210 \text{ km s}^{-1}$ measured by Smith *et al.* (1991). The line center shift between the broad and narrow component lines is $29 \pm 18 \text{ km s}^{-1}$, less than the uncertainty in broad component width, indicating a nearly edge-on viewing angle ($|90^\circ - \Theta| \leq 2^\circ$). The $\text{H}\alpha$ broad-to-narrow ratio is $0.84^{+0.03}_{-0.01}$, lying slightly above the extreme upper limit quoted by Smith *et al.* (1991), 0.73 ± 0.06 . The $\text{H}\beta$ broad-to-narrow ratio is somewhat larger, $0.93^{+0.18}_{-0.16}$. The $\text{H}\beta$ broad-to-narrow ratio is larger than that of $\text{H}\alpha$ due to the lower conversion efficiency of Ly γ photons into $\text{H}\beta$ in the narrow component (c.f. Ghavamian *et al.* 2001). The broad-to-narrow ratio of the $\text{H}\gamma$ line ($0.63^{+0.53}_{-0.25}$) is poorly constrained due to the faintness of the broad component. The observed $\text{H}\alpha$ surface brightness of the Balmer filament in SN 1006 is $(3.2 \pm 0.7) \times 10^{-16} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ arcsec}^{-2}$ (including the 20% photometric uncertainty), in reasonable agreement with the value quoted by KWC87.

5. ANALYSIS AND MODELS

To model the Balmer line profiles, we used the one-dimensional, plane-parallel shock code described in Ghavamian *et al.* (2001) to compute a grid of numerical models. Briefly, the code predicts the density and temperature of electrons, protons and hot neutrals behind a non-radiative shock. Broad to narrow flux ratios are computed by a Monte Carlo simulation which follows collisionally excited Ly β and Ly γ photons until they are absorbed or escape from the grid. Collisional excitation and ionization of H by both electrons and protons is included in the models. Free parameters of the model are the preshock neutral fraction f_{H^0} , shock speed, v_s , and fractional equilibration immediately behind the shock. The latter quantity denotes the degree of electron-ion equilibration by plasma turbulence at the shock front, where $f_{eq} = 0$ represents no (pure Coulomb) equilibration, while $f_{eq} = 1$ represents full equilibration. The electron and proton temperatures are then defined by the relation

$$T_e = \frac{3}{16} \frac{m_p v_s^2}{k} (\mu f_{eq} + \frac{m_e}{m_p} (1 - f_{eq})) \quad (1)$$

$$T_p = \frac{3}{16} \frac{m_p v_s^2}{k} (\mu f_{eq} + (1 - f_{eq})) \quad (2)$$

⁶IRAF is distributed by the National Optical Astronomy Observatories, which is operated by the AURA, Inc. under cooperative agreement with the National Science Foundation

where μ is the mean molecular weight ($=0.6$ for Galactic abundances, Allen 1973) and $\frac{3}{16}\mu m_p v_S^2$ is the mean postshock temperature. For an input f_{eq} , the subsequent evolution of T_e and T_p by Coulomb collisions is followed as a function of position behind the shock. We have updated the code to include the calculation of both the He ionization balance and the contribution of He ionization to the postshock electron density. To increase the signal-to-noise of our simulations, we have also increased the number of excitations followed by the Monte Carlo simulation from 10,000 used in the earlier paper (Ghavamian *et al.* 2001) to 50,000.

Using the same method outlined in Ghavamian *et al.* (2001), we computed a grid of models in equilibrium and shock velocity, combining each f_{eq} with the v_S required to match the observed FWHM of the broad component. We sampled the subset $f_{eq} = (0, 0.03, 0.05, 0.1, 0.2, 0.4, 0.6, 0.8, 1)$, corresponding to $v_S = (2865, 2880, 2895, 2920, 2965, 3100, 3235, 3395, 3580)$ km s $^{-1}$, respectively. The H α and H β broad-to-narrow ratios predicted by our models are shown in Fig. 5. It is immediately obvious that multiple combinations of f_{eq} and I_b/I_n are consistent with the observed broad-to-narrow ratios. This makes the unambiguous, simultaneous determination of f_{H^0} and f_{eq} from Fig. 5 impossible. In an earlier calculation, Laming (2000) obtained similar shapes for the broad-to-narrow ratio curves, the main difference being that our curves flatten out and then decline slightly at high equilibrations. The reason for the disagreement is that unlike our models, the Laming (2000) calculation maintains a constant shock speed while varying f_{eq} . The slight downturn in our computed narrow ratio curves is caused by the decline of the $H^+ - H^0$ charge exchange rate at high equilibrations (i.e., high shock speeds). As a result, the hot neutral density peaks at progressively smaller values behind the shock as v_S increases beyond 3000 km s $^{-1}$.

Fortunately, we can use the He I/He II and He I/H α flux ratios to break the degeneracy in Fig. 5. Since every He atom in the postshock ionization zone will spend some time as He^+ (the preshock He^{++} fraction typically amounts to only a few percent), the predicted He II $\lambda 4686$ flux is *insensitive* to the preshock He ionization fraction. Furthermore, as demonstrated by Hartigan (1999) and Laming *et al.* (1996), the He II $\lambda 4686$ flux is also insensitive to the degree of electron-ion temperature equilibration behind the shock (a result that is confirmed by our models). Therefore, we can use the He I/He II ratio to constrain the preshock He^0 fraction. Once this quantity is known, we can then use the He I/H α ratio to constrain the preshock H^0 fraction. This will enable us to determine the appropriate choice of broad-to-narrow ratio from Fig. 5. From our shock models we estimate that regardless of the assumed equilibration, 90% of the He II emission is produced within a zone $<1''$ thick (assuming a total preshock density of 1 cm $^{-3}$ (Winkler & Long 1997), shock speed 3000 km s $^{-1}$ and a distance of 2 kpc). This is smaller than both the pixel scale ($0''.51$ pixel $^{-1}$) and the slit width ($2''.0$). Therefore, we believe that all of the He II emission layer was contained within the slit.

In most interstellar shocks where He I emission is observed (mainly radiative shocks), the $\lambda 5876$ line is brighter than the $\lambda 6678$ line. The $\lambda 6678$ emission line arises from

the singlet transition $^1P(1s2p) - ^1D(1s3d)$ and is most efficiently excited from the ground state (i.e., by collisional excitation). The $\lambda 5876$ line, however, arises from a triplet transition and is most efficiently excited by cascades from high lying energy levels (i.e., from recombination). Therefore, since recombination is negligible in SN 1006 shocks, we expect the former line to be significantly brighter than the latter (indeed, $\lambda 5876$ emission is not detected in our optical spectrum). In the following analysis we limit our calculations to the $\lambda 6678$ line intensities.

To estimate the He I/H α ratio for SN 1006, we computed the total H α and He I $\lambda 6678$ fluxes through the shock front using the combinations of f_{eq} and v_S listed above. Calculation of the He I $\lambda 6678$ flux in SN 1006 is relatively straightforward because proton excitation and ionization of He is negligible (Laming *et al.* 1996). In addition, the only significant contribution to the He I flux other than $^1S(1s^2) - ^1D(1s3d)$ excitation are cascades from the 4f level via $^1D(1s3d) - ^1F(1s4f)$. We have included both direct excitation and cascade contributions to 3d in our He I $\lambda 6678$ flux calculation. We obtained the electron excitation cross sections from the convergent close coupling (CCC) results of Fursa & Bray (1995). At high energies, we approximated the cross sections with a Born-like dependence, scaled to match the CCC values at 0.9 keV. Our He ionization balance calculation uses the electron ionization rates of Younger (1981) for He^0 and He^+ . We also utilized the CCC cross sections of Fursa and Bray (1995) to compute the collisional excitation rate of He II $\lambda 4686$ electrons (proton excitation is negligible, see Laming *et al.* 1996). Given the low density, the order of magnitude lower abundance of He relative to H, and the large line width of He II, we assumed Case A conditions in the He II line, i.e., the radiative excitation of He II $\lambda 4686$ by He II $\lambda 243$ (Ly γ) absorption is negligible behind the shock.

The model F(6678)/F(4686) ratios are shown in Fig. 6 for a range of preshock He^0 fractions and temperature equilibrations. The dashed lines indicate error bars on the observed line ratio after dereddening. Since the ratio of the He I excitation rate to the He I ionization rate depends steeply on electron temperature for $T_e \lesssim 10^5$ K, the He I/He II ratio declines sharply at low values of f_{eq} (≤ 0.01). It is clear from this figure alone that matching the predicted He I/He II ratio with the observed value requires a preshock He neutral fraction ≥ 0.5 and low electron-ion equilibration.

The predicted He I/H α ratios are displayed in Fig. 7 for a range of preshock H neutral fractions. For simplicity, we assumed $f_{He^0} = 1$ in these models (varying this quantity merely moves the curves up or down). To facilitate comparison with the broad-to-narrow ratios, we used the same set of neutral fractions and combinations (f_{eq} , v_S) employed by the models shown in Fig. 5.

(1) Case A in both the broad and narrow components, and (2) Case A in the broad component, Case B in the narrow component. Since the inferred Ly β and Ly γ optical depths behind Balmer-dominated shocks are often $\lesssim 1$ in the narrow component line (Chevalier, Kirshner & Raymond 1980, Ghavamian 1999, Ghavamian *et al.* 2001), the Case A and Case B line ratio curves mark the largest and smallest values of He I/H α attainable for each combination

of preshock H and He ionization fraction.

Comparing Figs. 6 and 7, we find that (1) the observed He I/He II and He I/H α ratios are simultaneously matched for $f_{H^0} \sim 0.1$ and $f_{He^0} \geq 0.7$ in the preshock gas, and (2) only the lowest equilibration models ($f_{eq} \leq 0.05$) can match the observations. Returning to Fig. 5 and choosing the I_b/I_n curves for $f_{H^0} = 0.1$, we find a match between the observed and modeled H α and H β broad-to-narrow ratios for $f_{eq} \leq 0.1$, or $T_e/T_p \leq 0.07$. This corresponds to a shock speed of $2890 \pm 100 \text{ km s}^{-1}$ (including the measurement uncertainty of the broad component). This result is in good agreement with the shock speed estimated by Laming *et al.* (1996) from ultraviolet line analysis of HUT data: $2600 \pm 300 \text{ km s}^{-1}$. It also agrees with the low equilibration inferred by Laming *et al.* (1996): $T_e/T_i \leq 0.05$.

Given the intrinsic faintness of the He II emission, it is useful to compare the measured He II $\lambda 4686$ surface brightness with the value inferred from observation of the He II $\lambda 1640$ line by HUT (Raymond *et al.* 1995). Assuming Case A conditions in the He II lines (reasonable due to the low density of He), the predicted He II $\lambda 4686$ flux in the RC Spectrograph slit is

$$F(4686) = \frac{\nu_{4686}}{\nu_{1640}} \frac{\gamma_{4686}}{\gamma_{1640}} \frac{F_{RCSpec}(H\alpha)}{F_{HUT}(H\alpha)} F_{HUT}(1640) \quad (3)$$

where γ_{4686} is the number of $\lambda 4686$ photons per incident He⁺ ion and γ_{1640} is the corresponding quantity for the $\lambda 1640$ line. We have used our H α image (the star-subtracted version of Fig. 1) to scale from the He II $\lambda 4686$ flux expected in the HUT aperture to the value expected in the RC Spectrograph slit. This scaling is valid if the H neutral fractions are roughly equal in the two filaments (reasonable given their close proximity and comparable surface brightnesses). From the H α image we estimate $F_{RCSpec}(H\alpha)/F_{HUT}(H\alpha) \approx 0.13$. Next, from the calculations of Hartigan (2000) we estimate $\gamma_{4686}/\gamma_{1640} \approx 0.1$ for $10^5 \text{ K} \leq T_e \leq 10^7 \text{ K}$. Inserting these values into Eqn. 3 and using $F_{HUT}(1640)$ from Raymond *et al.* (1995, who also assumed $E(B - V) = 0.11$), we find $F(4686) = (1.1 \pm 0.2) \times 10^{-15} \text{ ergs cm}^{-2} \text{ s}^{-1}$. This agrees well with the $\lambda 4686$ flux we derive from our data, $(9.4 \pm 6.4) \times 10^{-16} \text{ ergs cm}^{-2} \text{ s}^{-1}$.

6. DISCUSSION

In their detailed imaging study of SN 1006, Winkler & Long (1997) found that X-ray emission in the NW peaks approximately $12''$ behind the optical filaments. Using nonequilibrium ionization models convolved with the ROSAT HRI response function, they found that a total preshock density $\sim 1 \text{ cm}^{-3}$ was required to match the spatial distribution of the X-ray emission. Combined with our estimated H preshock neutral fraction, this implies a preshock neutral density $n_{H^0} \sim 0.1 \text{ cm}^{-3}$ with $\tau_{Ly\beta} \sim 0.5$ in the narrow component. The inferred value of $\tau_{Ly\beta}$ will be even lower if Winkler & Long (1997) overestimated the total preshock density. The dereddened narrow Balmer decrement is 3.35 ± 0.14 , significantly lower than the value of 4.3–6 observed in Tycho's SNR (Ghavamian 2001). The absorption of narrow component Lyman photons within the shock structure scales linearly with f_{H^0} , as long as f_{H^0} is small ($\lesssim 0.5$). This enhances the H α emission more than H β , steepening the narrow component Balmer decrement.

Therefore, the modest decrement we observe in SN 1006 is consistent with our low derived Lyman optical depth.

6.1. The Preshock Ionization Fractions

The detection of the He I $\lambda 6678$ line indicates that a significant fraction of the preshock He is neutral. The collisional excitation of He II $\lambda 304$ behind the shock certainly produces some photoionization and heating of the ambient gas (see, for example, Tycho's SNR, Ghavamian *et al.* 2000). Since the photoionization rate of He⁰ by $\lambda 304$ photons is nearly an order of magnitude higher than the photoionization of H⁰ (Ghavamian *et al.* 2000), some additional mechanism is required which can ionize ambient H⁰ more effectively than He⁰. There are several possibilities which we discuss in turn below.

6.1.1. Photoionization From the Reverse Shock?

One process which may generate high H ionization fractions in the ambient medium while leaving most He neutral is photoionization by O II–O VIII emission from the outermost layer of shocked ejecta (e.g., Hamilton & Fesen 1988, hereafter HF88). The strongest of these lines have energies lying between the H and He⁰ ionization thresholds. Using the calculations of HF88 and the He II $\lambda 304$ photoionization models of Ghavamian *et al.* (2000), we can estimate the ratio of the H photoionization rate to the He photoionization rate in the interstellar medium due to radiation from the blast wave and shocked ejecta. Since each O atom entering the reverse shock will pass through all ionization stages between O⁺⁺ and O⁺⁷, the ratio of photoionizations $\alpha_{H/He}(R)$ at a distance R from the remnant center is

$$\alpha_{H/He}(R) \approx \frac{\sum_i \gamma(\lambda_i) \sigma^H(\lambda_i) + \eta \gamma(304) \sigma^H(304)}{\sum_i \gamma(\lambda_i) \sigma^{He}(\lambda_i) + \eta \gamma(304) \sigma^{He}(304)} \quad (4)$$

where

$$\eta = \frac{\frac{N_{bw}(He)}{(R - R_{bw})^2}}{\frac{N_{ej}(O)}{(R - R_{ej}(O))^2}}$$

Here, $\gamma(\lambda_i)$ is the number of ionizing photons per ion of wavelength λ_i (the Bethe parameter as tabulated by HF88 for the reverse shock), $\gamma(304)$ is the Bethe parameter for He II $\lambda 304$ photons from the blast wave (≈ 1.1 for $v_S > 2000 \text{ km s}^{-1}$, Ghavamian *et al.* 2000), σ^H is the H⁰ photoionization cross section from Osterbrock (1989) and σ^{He} is the corresponding cross section for He⁰ (Reilman & Manson 1979). The current O-layer and blast wave radii are $R_{ej}(O)$ and R_{bw} , respectively, while $N_{ej}(O)$ is the number of shocked O atoms in the ejecta and $N_{bw}(He)$ is the number of He atoms swept up by the blast wave. In our calculation we have assumed that the ionization state of the ISM is negligibly affected by recombination. HF88 found that the relative masses of Fe, Si and O predicted by carbon deflagration model CDTG7 (Woosley & Weaver 1987) matched the observed Fe II $\lambda\lambda 2586, 2599$ absorption profile in SN 1006 best. Therefore, using the shocked O ejecta mass of $0.41 M_\odot$ and blast wave mass of $4.3 M_\odot$ from model CDTG7, $R_{ej}(O) = 6 \text{ pc}$ and $R_{bw} = 8.3 \text{ pc}$ from HF88 and neglecting radiative transfer effects on the

photoionizing radiation, we find $\alpha_{H/He} \sim 0.1$ for a parcel of interstellar gas just outside the blast wave radius. This result suggests that reverse shock photoionization cannot produce our derived low H neutral fraction. The smallness of $\alpha_{H/He}$ is mainly due to that fact that (1) there are nearly an order of magnitude more He atoms in the blast wave than there are O atoms in the reverse shock, and (2) since $R_{ej}(O) < R_{bw}$, a given parcel of interstellar gas sees a larger emitting area from the blast wave than from the reverse shock.

A more detailed, time dependent calculation would be desirable. However, we note that the reverse shock has already propagated through the entire O-layer, while the blast wave continues to accumulate interstellar material (and hence He atoms). Therefore, the ratio $M_{ej}(O)/M_{bw}$ is decreasing as the supernova remnant evolves and the blast wave is becoming the increasingly dominant photoionization source for the ISM around SN 1006. Finally, we note that even if we had included photoionization from the Si-rich and Fe-rich layers in Eqn. 4, our conclusions would remain unchanged since most of the EUV/X-ray lines from Si and Fe will ionize He much more effectively than H.

6.1.2. Relic Ionization from the Supernova Explosion?

An interesting source of photoionization is the EUV flash produced when photons escape from the detonating white dwarf (Type Ia SN). Using the EUV burst models presented by Blinnikov & Sorokina (2000) and Sorokina & Blinnikov (2001, in preparation), we estimate $\sim 10^{55} - 10^{56}$ ionizing photons ($E > 13.6$ eV) are released over the integrated Type Ia SN light curve. The precise number can vary by a factor ~ 5 , depending on which detonation model is used in the light curve calculation (see Blinnikov & Sorokina (2000) for model descriptions). Due to the steepness of the EUV spectrum, the H^0 ionization rate is much larger than the He^0 ionization rate, a desirable feature. However, assuming an average ISM density of $0.06 - 1.0$ cm^{-3} around SN 1006 and distance of 2 kpc, there were $\gtrsim 5 \times 10^{57}$ H atoms contained within the current blast wave radius at the time of the SN explosion. Even if these H atoms were already partially ionized beforehand, the total ionizing flux from a Type Ia explosion still falls an order of magnitude short of the value needed to effectively ionize the gas at the current blast wave radius. The disparity is worsened if we include radiative transfer loss of the released photons as they (1) propagate through the ISM, and (2) possibly encounter a circumstellar envelope left by a nondegenerate donor star prior to the explosion (Cumming *et al.* 1996). Therefore, regardless of whether the EUV emission ionizes H more effectively than He, there are not enough photons produced in the explosion to explain our inferred neutral fractions. This result also agrees with the SN flash ionization calculations of Cheng & Bruhweiler (1996).

6.1.3. Photoionization from Galactic Background Radiation?

If neither the supernova event nor the postshock emission from the supernova remnant can account for our deduced H^0 and He^0 fractions, we must consider nonlocal sources of ionizing radiation. Given the lack of late O and

early B stars in the vicinity of SN 1006, the most promising source is the Galactic ionizing background. Reynolds (1984) observed the $H\alpha$ recombination line toward a sample of high Galactic latitude pulsars ($|b| > 5$) at distances of 2–3 kpc. He found that Lyman continuum radiation from O stars in the Galactic plane was the most likely source of ionization in the diffuse ISM. Since the Lyman continuum is produced by recombination, it declines nearly exponentially beyond each ionization edge (Brussard & Van De Hulst 1962). Therefore, incident Lyman continuum photons will ionize H far more effectively than He, a desirable property.

The remaining question is whether the Galactic Lyman continuum flux is high enough at the position of SN 1006 ($b = +14.6^\circ$) to produce the necessary amount of H ionization. Assuming a density of 1 cm^{-3} around SN 1006, the recombination rate for $5000 \text{ K} \lesssim T \lesssim 10,000 \text{ K}$ is $\sim 5 \times 10^{-13} \text{ s}^{-1}$ (Osterbrock 1989). Therefore, obtaining an H neutral fraction ~ 0.1 requires a photoionization rate $\sim 5 \times 10^{-12} \text{ s}^{-1}$. Dove, Shull & Ferrara (2000) estimated a Lyman continuum flux $\sim 5 \times 10^7 \text{ photons cm}^{-2} \text{ s}^{-1}$ from O stars in the Galactic disk. In their time dependent radiative transfer calculation for the ionizing continuum, they estimated that roughly 3%–6% of the Lyman continuum photons escape through superbubbles and Galactic chimneys into the halo of the Milky Way, producing the “Reynolds Layer” of diffuse $H\alpha$ emission (Reynolds 1984). If the Lyman photons propagating into the halo suffer no further attenuation, the corresponding H ionization rate in the vicinity of SN 1006 is approximately $F_{Ly} \sigma^H(912) \sim 1 - 2 \times 10^{-11} \text{ photons s}^{-1}$, more than enough to produce the necessary H ionization around SN 1006.

If the Galactic Lyman continuum background preionizes most of the ambient H on the NW side of SN 1006, the question arises of why the ambient H on the eastern half of Tycho’s SNR (also the remnant of a Type Ia explosion) was found to be mostly neutral (Ghavamian *et al.* 2000). First, the eastern half of Tycho’s SNR is likely interacting with the outer edge of a warm H I cloud (Reynoso *et al.* 1999, Ghavamian *et al.* 2000). Second, the presence of large structures in the Galactic disk (H I clouds, superbubbles, molecular cloud complexes, etc.) produces large variations in the Lyman continuum flux from position to position and time to time (Dove, Shull & Ferrara 2000). Since Tycho’s SNR is close to the Galactic plane ($b = +1.4^\circ$), it is possible that this remnant is shielded from the Lyman continuum emission by the many dense clouds in the Galactic disk. SN 1006, on the other hand, lies nearly 500 pc above the Galactic plane. Therefore, neutral H near SN 1006 may “see” a larger time averaged Lyman continuum than the neutral H around Tycho’s SNR.

6.2. Comparison With Ionization Estimates of the Local Bubble

An interesting comparison can be drawn between our derived H and He neutral fractions and values estimated for the local ($\lesssim 200$ pc) ISM. For example, using EUVE spectra of 29 DA white dwarfs, Wolff *et al.* (1999) estimated $f_{He^+} \sim 0.4$ in the Local Bubble, with no evidence of a gradient or other general pattern. They interpreted this result as evidence for a lack of He-ionizing sources in any partic-

ular location. On the other hand, Wolff *et al.* (1999) found a positive gradient in the H ionization fraction toward the Canis Major cavity, implicating the B-stars Adara and Mirzam as the photoionizing sources. In some respects the ionization properties of the Local Bubble and the ISM near SN 1006 are similar: the H ionization is strongly influenced by EUV radiation, while He remains largely unaffected. The main difference between the two cases, of course, is the lack of a local H-ionizing source near SN 1006. Our derived He ionization fraction (≤ 0.3) for SN 1006 is lower than the average value of 0.4 ± 0.1 measured by Wolff *et al.* (1999), but there are individual cases from their WD sample with He ionization fractions similar to ours.

6.3. Comparison With Earlier Results

Our optical analysis lends strong support to the picture developed by earlier X-ray, UV and optical studies of SN 1006. Multi-epoch optical imagery of SN 1006 has yielded a proper motion of $0.3'' \pm 0.04'' \text{ yr}^{-1}$ (Long, Blair & van den Bergh 1988) for the Balmer filaments in the NW. Combined with our estimated shock speed this gives a distance of $2.0^{+0.4}_{-0.3} \text{ kpc}$, agreeing well with the value 1.8 ± 0.3 obtained by Laming *et al.* (1996). The SN 1006 blast wave is the fastest shock we have thus far modeled with our Balmer-dominated shock code. In modeling Balmer spectra from the Cygnus Loop ($v_S \sim 300 \text{ km s}^{-1}$), RCW 86 ($v_S \sim 600 \text{ km s}^{-1}$) and Tycho's SNR ($v_S \sim 2000 \text{ km s}^{-1}$), we noted that the equilibration f_{eq} declined from near 1 for the slowest shocks to $\lesssim 0.2$ for the fastest shocks (Ghavamian 1999, Ghavamian *et al.* 2001). The shocks in SN 1006 are closer to 3000 km s^{-1} and are matched by an even lower f_{eq} , further evidence of an inverse correlation between shock speed and equilibration. The case for low equilibration at high Mach numbers has recently received further support from X-ray observations of SN 1987A. In their analysis of high resolution HETG data from *Chandra*, Michael *et al.* (2001) found that a high blast wave speed (3500 km s^{-1}) along with a low electron temperature ($f_{eq} < 0.1$) were required to match the X-ray profiles of the SN 1987A spectrum.

6.4. Implications for Collisionless Shock Physics

There are few numerical simulations available of high Mach number, collisionless shocks. Cargill & Papadopoulos (1988) modeled the electron-ion heating in a shock moving perpendicular to the ambient magnetic field, with $M_A = 50$ ($M_A \equiv v_S/(c_s^2 + v_A^2)^{1/2}$ is the Alfvénic Mach number, where c_s is the sound speed and v_A is the Alfvén speed of the preshock gas). In their hybrid simulation the electrons were treated as a resistive fluid, while the ions were treated as individual particles. Cargill & Papadopoulos (1988) found that two-stream plasma instabilities produced by ions reflected upstream resulted in $T_e/T_i \sim 0.25$ at the shock front. Shimada & Hoshino (2000) tried an approach where both electrons and ions were treated as particles. At the highest Mach numbers covered by their simulations, they predicted that the equilibration would increase with Mach number. In their fastest simulated shock ($M_A = 20$), they found $T_e/T_i \sim 0.2$. Their curve of T_e/T_i vs. M_A flattens considerably as $M_A \rightarrow 20$; extrapolating this curve to $M_A \gtrsim 200$ leads to significantly

larger equilibration than implied by the ultraviolet (Laming *et al.* 1996) and optical (this paper) observations of SN 1006.

During the last few years, a considerable effort has been made in modeling collisionless heating processes in solar wind shocks ($M_A \leq 20$). Hull & Scudder (2000) combined an electron fluid treatment with constraints from ISEE 1 data of Earth's bow shock to model the electron-ion equilibration process. For the first time, they were able to match the observed electron temperature jumps in fast shocks propagating both parallel and perpendicular to the magnetic field. These models emphasize the role of coherent (DC) electric forces in the collisionless heating, rather than wave-particle interactions. The advantage of the Hull & Scudder (2000) treatment is that it draws upon in situ observations of real collisionless shocks rather than computer simulations. Both these models and the observations (Schwartz *et al.* 1988) predict $T_e/T_i \sim 0.25$ for shocks in the solar wind. Although we are unable to perform similar detailed observations of supernova shocks, extending the models of Hull & Scudder (2000) to much higher Mach numbers may be an interesting direction for future work on collisionless shocks.

7. SUMMARY

The optical spectrum presented here is the deepest yet obtained of the nonradiative filament along the northwest rim of SN 1006. Lines of H α , H β , H γ , H δ , He I $\lambda 6678$ and possibly He II $\lambda 4686$ have been detected—more lines than in the spectrum of any other purely nonradiative SNR filament. We have detected both broad and narrow component emission lines of H α , H β and H γ . The broad and narrow components result, respectively, from fast neutrals produced by charge exchange and slow neutrals that are collisionally excited.

The broad component widths, the ratios of broad to narrow flux in H α and H β , and the Balmer line strengths relative to He I and He II enable us to determine the important parameters for the shock. We find a shock speed of $2890 \pm 100 \text{ km s}^{-1}$ and a very low degree of electron-ion equilibration, $T_e/T_p \leq 0.07$, both consistent with earlier measurements. In addition, our models can match the observations only if the preshock H is mostly ionized ($\sim 90\%$) and preshock He is mostly neutral ($\gtrsim 70\%$). In seeking an explanation for this ionization structure, we have concluded that neither radiation from the blast wave nor radiation from the reverse shock can effectively ionize the ambient H surrounding SN 1006. We also find that the EUV flash from the supernova explosion is incapable of producing the necessary H ionization. The most plausible explanation for the high H ionization fraction seems to be photoionization from the Galactic Lyman continuum. Finally, comparing our derived electron-ion equilibration with values derived for other nonradiative SNRs reinforces the earlier conclusion that the fastest shocks have the least equilibration. A thorough understanding of this trend remains as a challenge for the study of collisionless shock physics.

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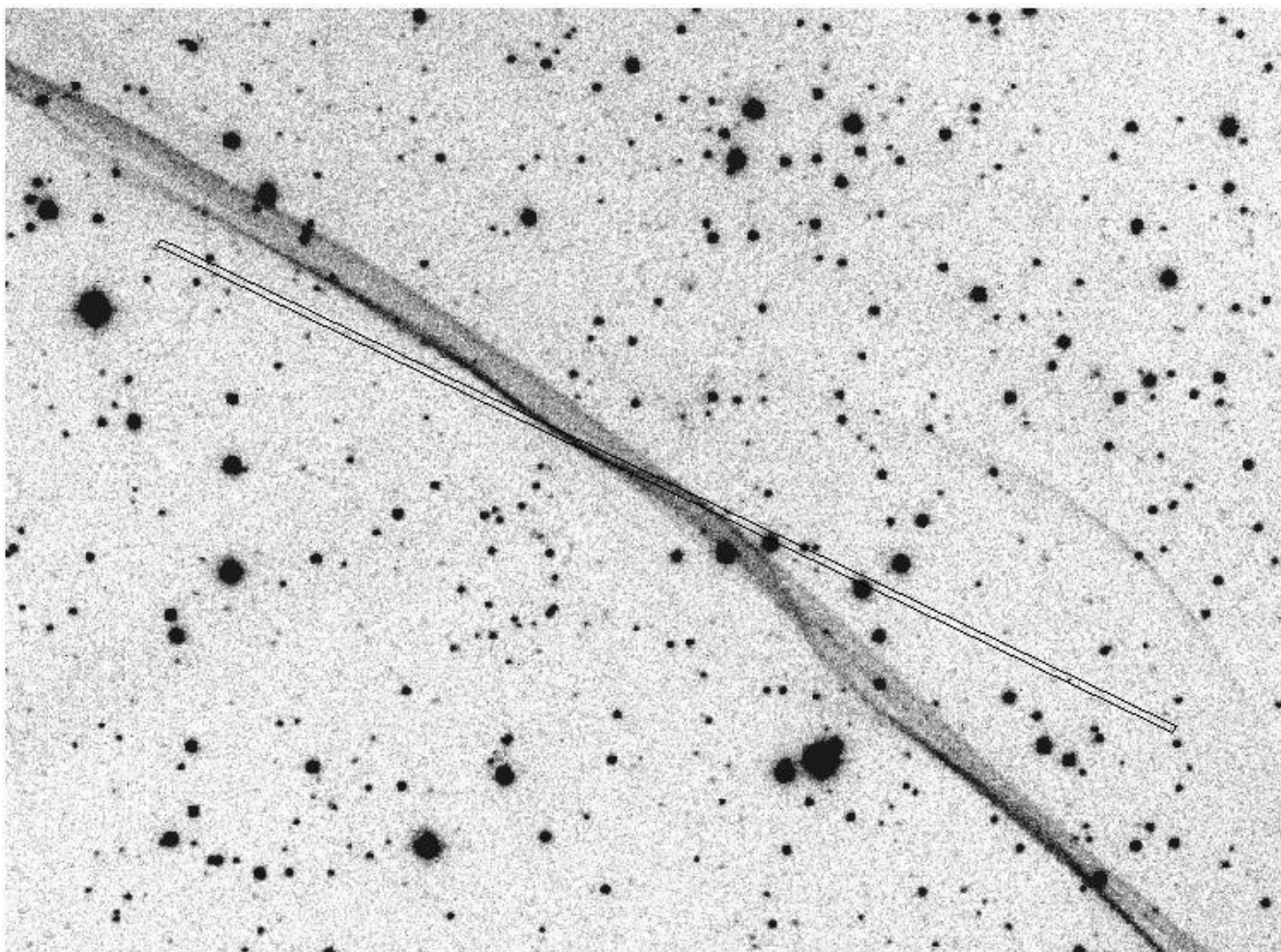


FIG. 1.— A narrow band $H\alpha$ image of NW SN 1006, acquired from the CTIO 0.9-m telescope in 1998 June, shortly before the spectroscopy reported here. The location and PA of the $2'' \times 5'$ RC Spectrograph slit is marked. North is at the top, East is to the left.

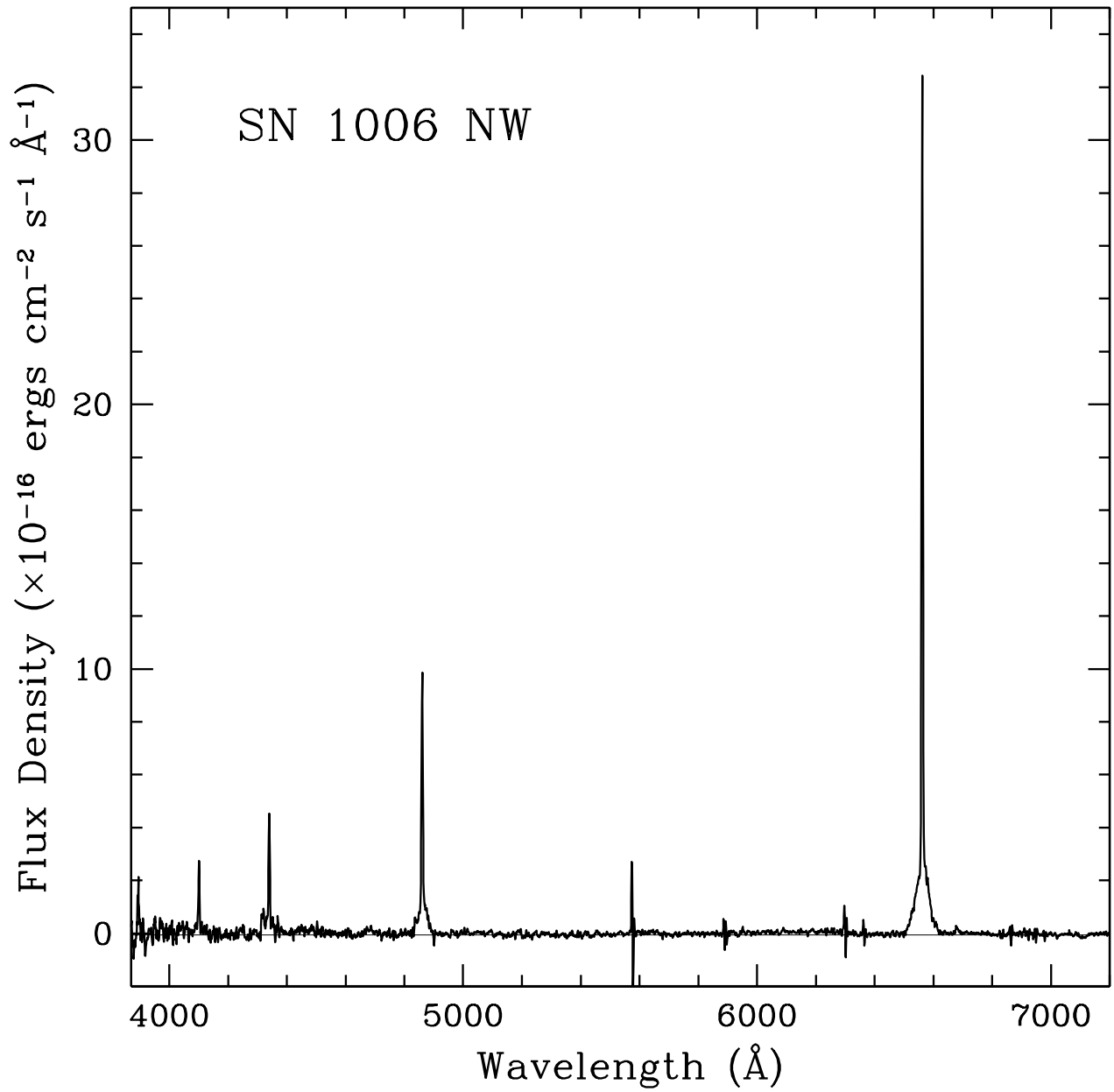


FIG. 2.— Full one-dimensional spectrum of the NW SN 1006 Balmer filament. Narrow emission lines of $H\alpha$, $H\beta$, $H\gamma$ and $H\delta$ are detected, along with broad emission lines in $H\alpha$, $H\beta$, and possibly $H\gamma$. The sharp feature near the middle of the spectrum is an artifact of the sky subtraction.

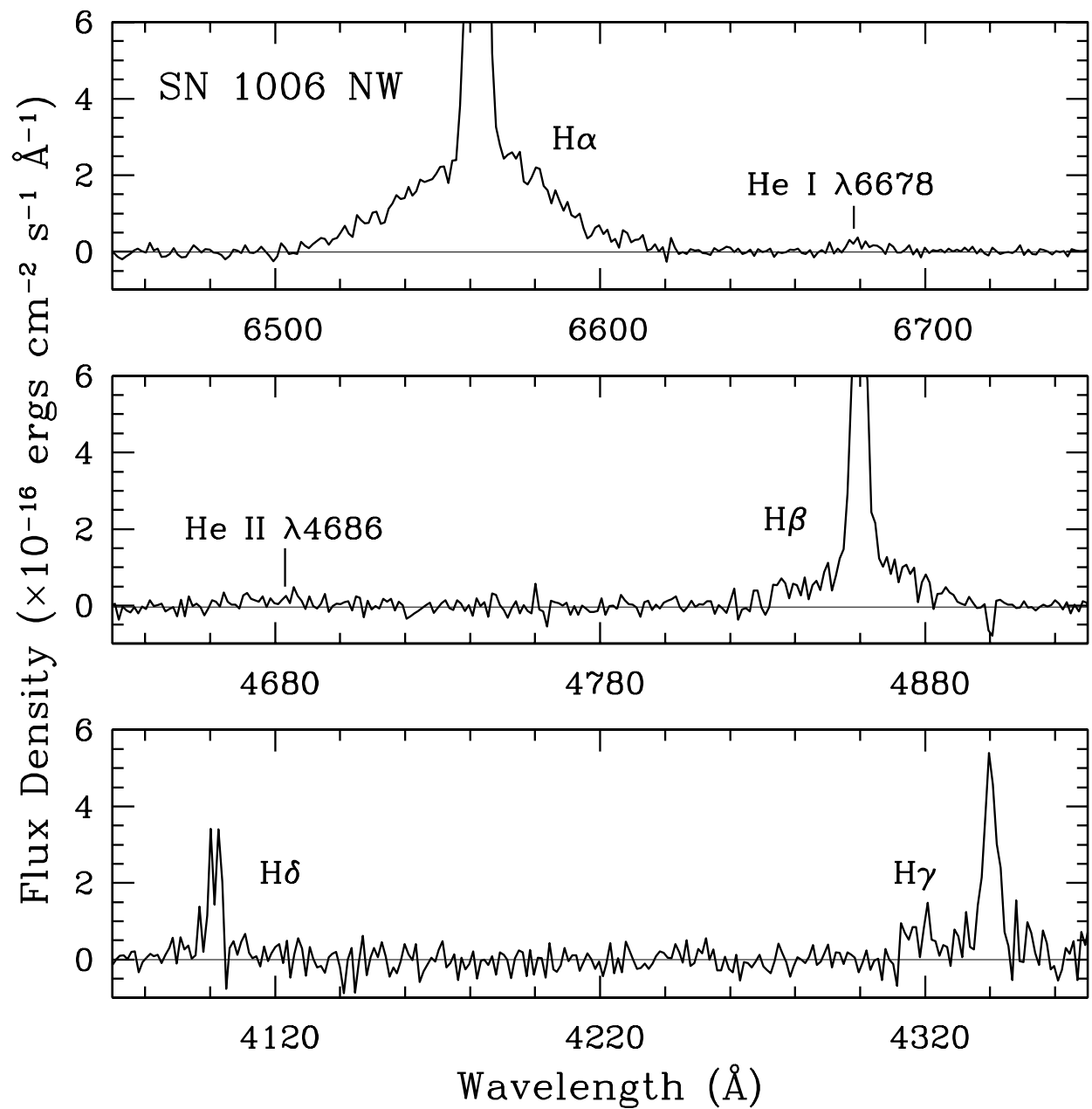


FIG. 3.— Closeup view of the NW SN 1006 Balmer spectrum. No smoothing has been applied. Among the newly detected lines is He I $\lambda 6678$. There may also be a weak detection of the He II $\lambda 4686$ line.

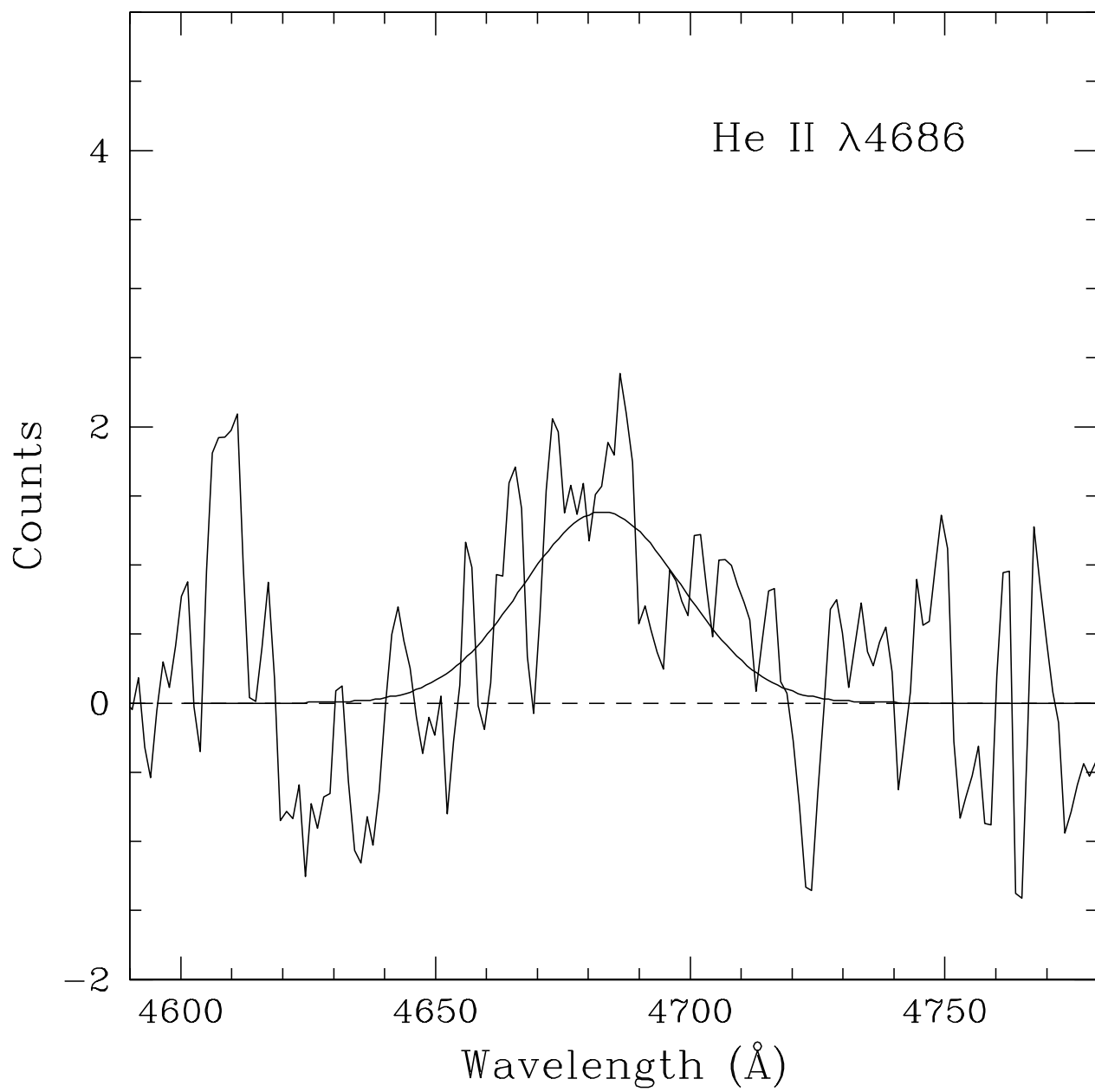


FIG. 4.— The He II $\lambda 4686$ line profile. The fit has been performed assuming the He II line width equals that of H α .

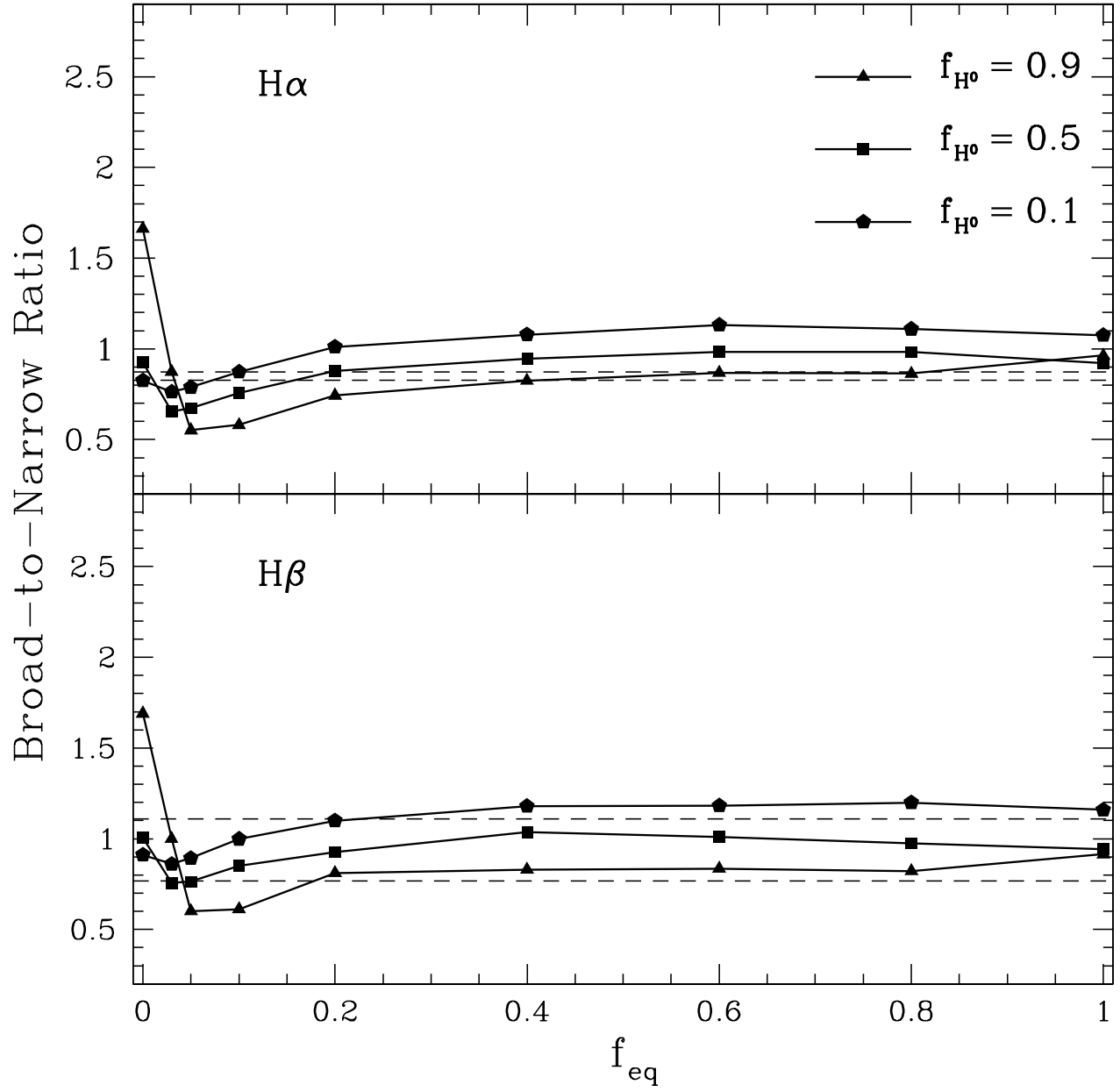


FIG. 5.— The calculated $H\alpha$ and $H\beta$ broad-to-narrow ratios vs. the electron-ion equilibration fraction f_{eq} at the shock front. The dashed horizontal lines mark the upper and lower limits of the observed broad-to-narrow ratios. Curves of constant preshock neutral fraction are shown for cases where the preshock H is mostly neutral, half neutral and mostly ionized.

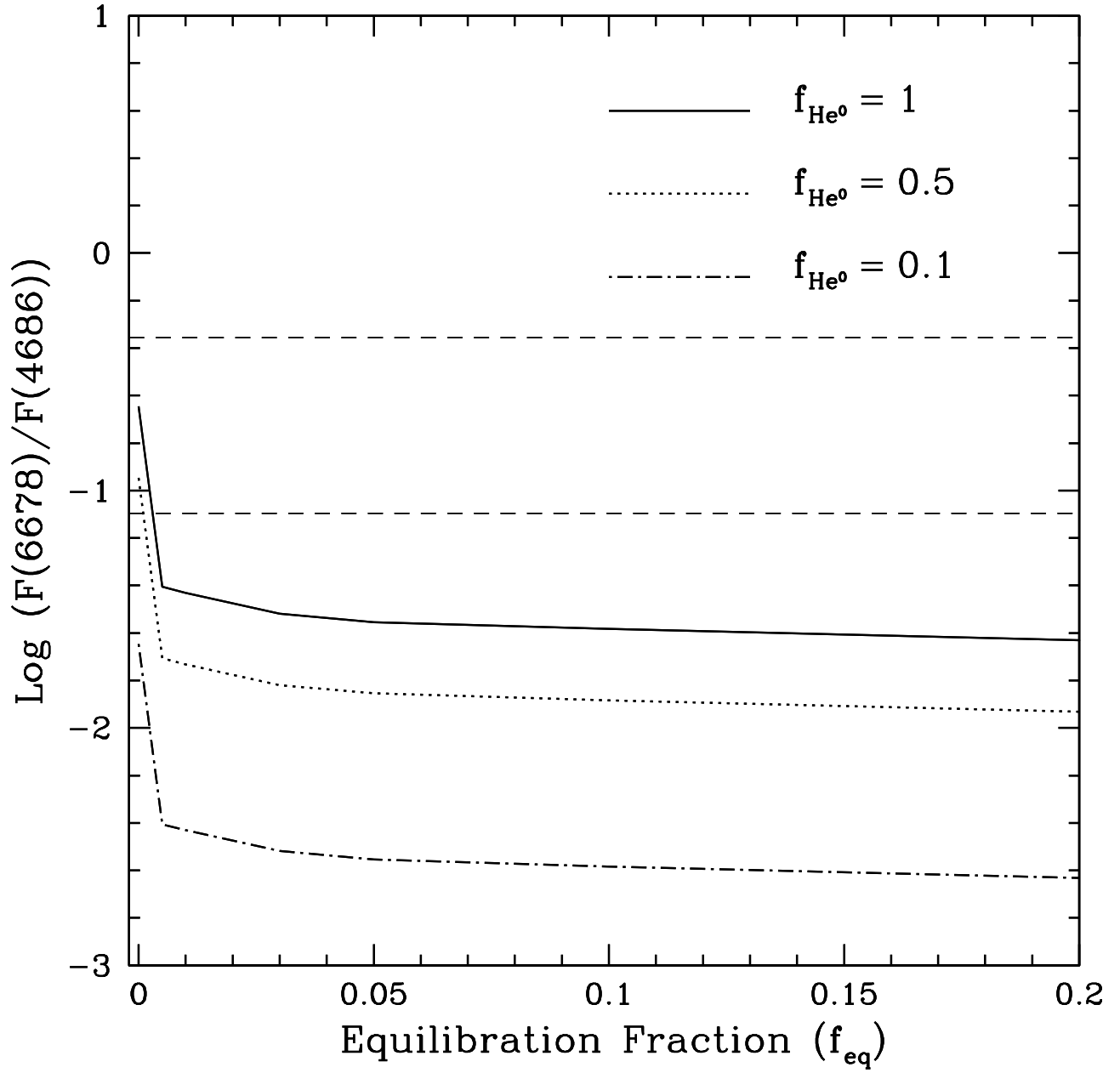


FIG. 6.— The observed vs. predicted He I/He II line ratio, shown for a range of preshock He neutral fractions. The observed range of the He I/He II ratio is indicated by the horizontal dashed lines. The model ratios have been computed assuming that neutral H in the narrow component is optically thin to Ly β photons behind the shock (Case A).

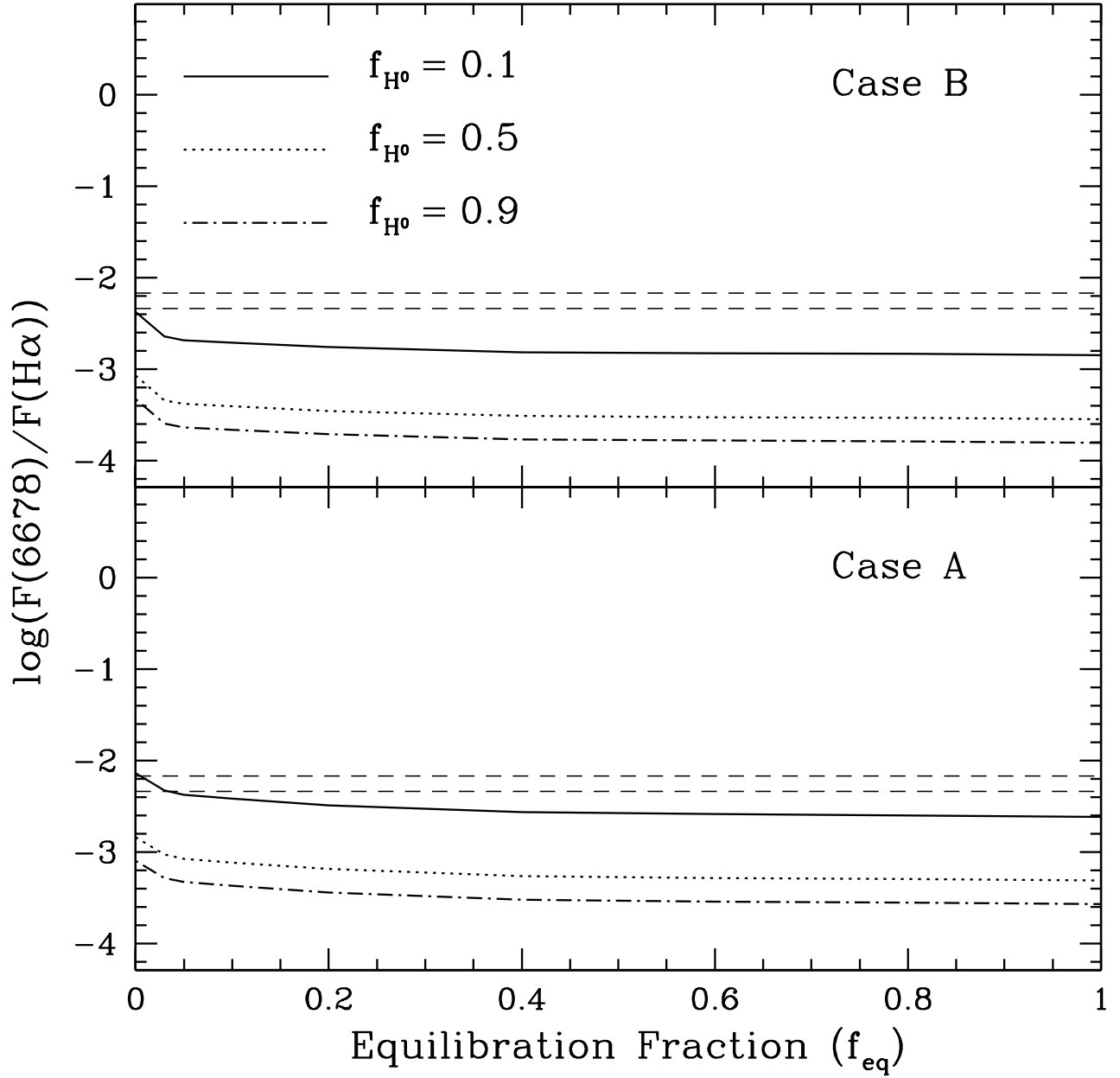


FIG. 7.— The dependence of the He I/ $H\alpha$ ratio on f_{eq} , shown for a sampling of preshock H neutral fractions. The preshock He is assumed to be fully neutral. The dashed lines indicate the observed range of the He I/ $H\alpha$ flux ratio. The ratios in the upper figure are computed assuming that neutral H in the narrow component is optically thick to Ly β behind the shock (Case B), while the ratios in the lower figure assume that neutral H in narrow component is optically thin to Ly β behind the shock (Case A).

TABLE 1
OBSERVED SURFACE BRIGHTNESSES FOR NW SN 1006^a

Line	$F(\lambda)_{obs}$	$I(\lambda)^b$	Broad to Narrow Ratio	Velocity Width (km s^{-1})
H α	355.3 ± 34.8	318.1 ± 31.2	$0.84^{+0.03}_{-0.01}$	2290 ± 80
H β	100 ^c	100 ^d	$0.93^{+0.18}_{-0.16}$	2290 (fixed)
H γ	36.4 ± 7.4	38.1 ± 7.7	$0.63^{+0.53}_{-0.25}$	2290 (fixed)
H δ	> 16.8	> 18.3	...	215 (fixed) ^e
He I $\lambda 6678$	2.0 ± 0.4	1.9 ± 0.44	— ^f	215 (fixed)
He II $\lambda 4686$	6.6 ± 4.5	7.0 ± 4.8	—	2290 (fixed)

^aBalmer line surface brightnesses are quoted for the combined broad and narrow component emission from a $2'' \times 51''$ section of filament. The error bars are 1σ .

^bDereddened intensities assume $E(B-V) = 0.11 \pm 0.02$ (Schweizer & Middleditch 1980).

^cObserved total H β surface brightness of $9.0 \times 10^{-17} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ arcsec}^{-2}$.

^dDereddened total H β surface brightness of $1.3 \times 10^{-16} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ arcsec}^{-2}$.

^eNarrow component line widths are fixed at 205 km s^{-1} (the instrumental resolution).

^fThe He I line is intrinsically narrow, while the He II line width is intrinsically broad.